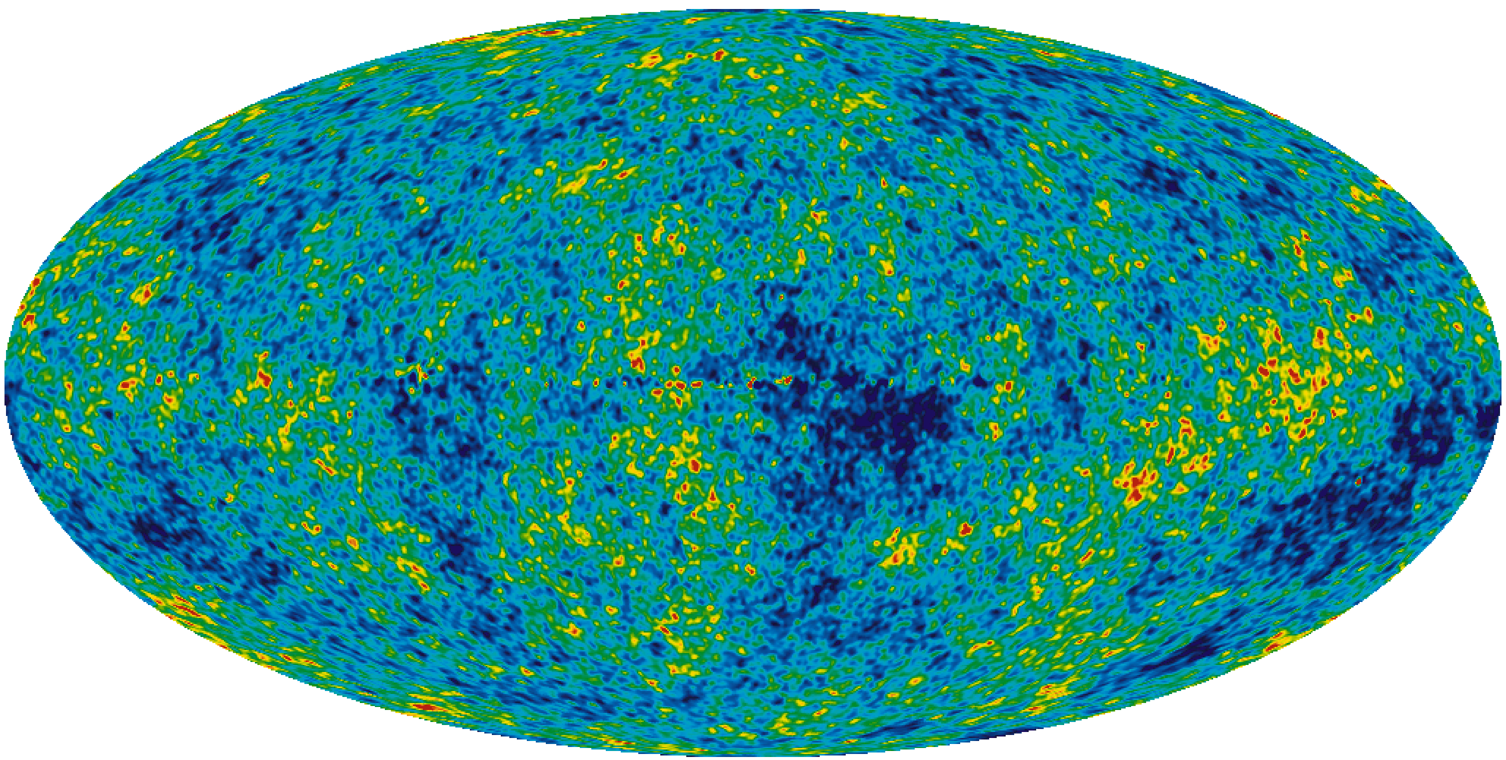
# Mysteries on universe's largest observable scales

Dragan Huterer

Kavli Institute for Cosmological Physics, Chicago

Collaborators:
Craig Copi (CWRU), Dominik Schwarz (Bielefeld), Glenn Starkman (CWRU)
Chris Gordon, Wayne Hu, Tom Crawford (Chicago)

# How does the universe look at largest observable scales?



ILC map, WMAP collaboration

# Outline

Motivation and overview of concurrent findings

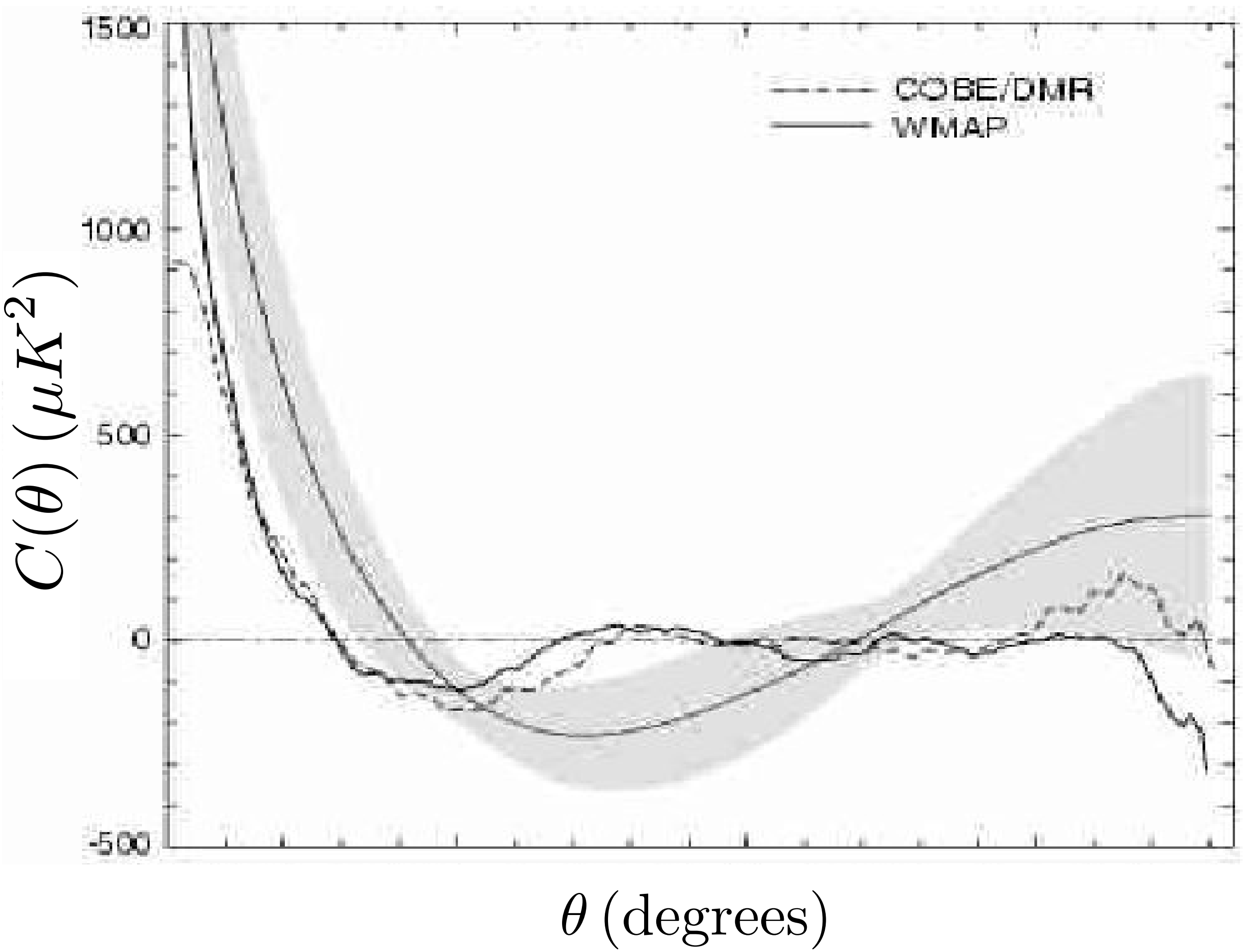
Multipole Vectors

Large-scale alignments

Various explanations

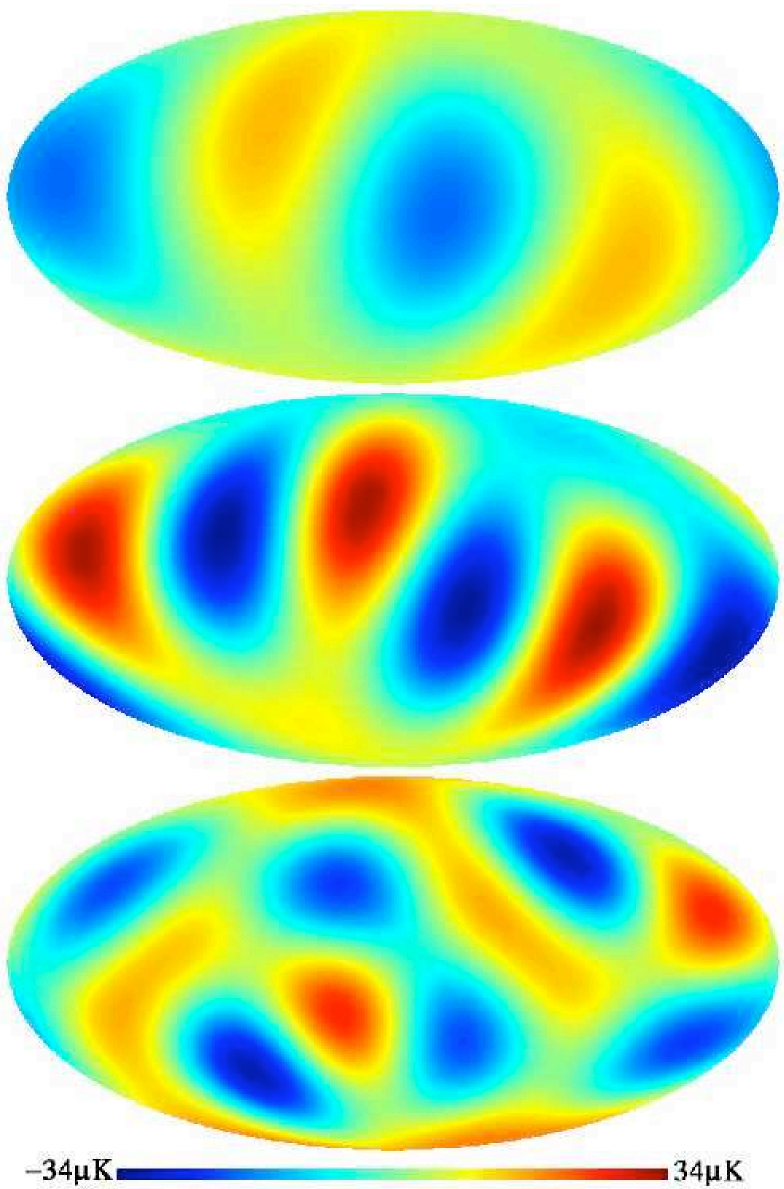
Future prospects and conclusions

# Low power on large scales




Spergel et al 2003: 0.2% of sims have less power at angles >60 deg

# l=2, 3 are aligned and planar



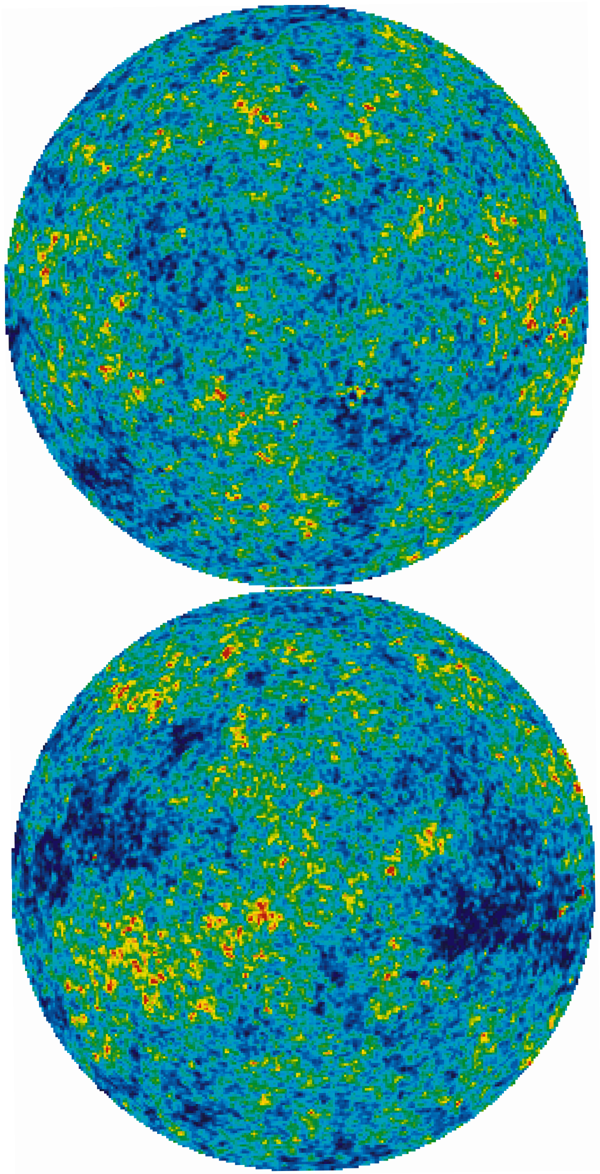

$$\hat{L}_\ell^2 \equiv \frac{\sum\limits_{m=-\ell}^{\ell} m^2 \left|a_{\ell m}\right|^2}{\ell^2 \sum\limits_{m=-\ell}^{\ell} \left|a_{\ell m}\right|^2}$$

l=3 is planar: P~1/20

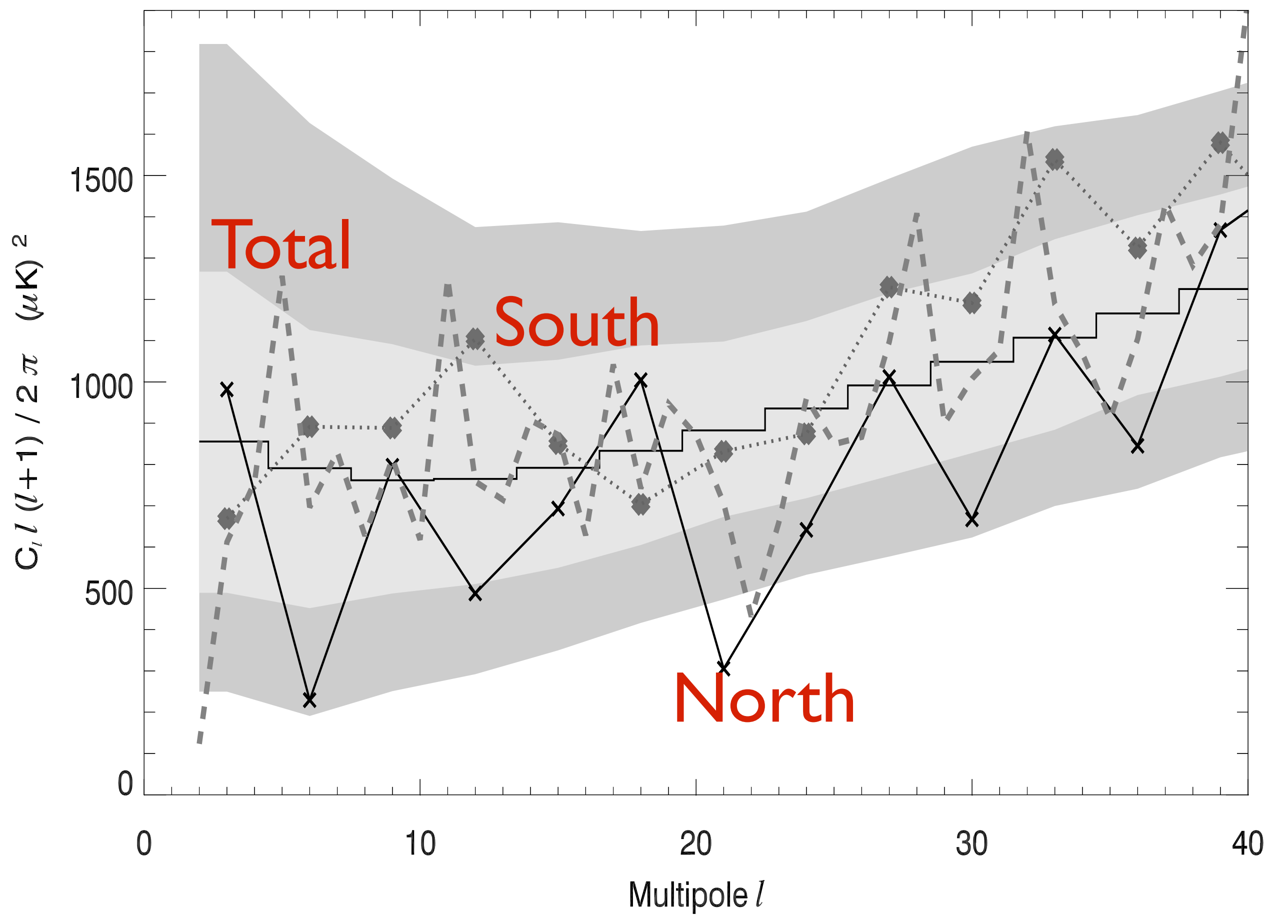
l=2,3 is are aligned: P~1/60

de Oliveira-Costa, Tegmark, Zaldarriaga & Hamilton 2004

# N/S power asymmetry



South (ecliptic) has more power than north




Eriksen et al 2004;
Hansen, Banday and Gorski 2004

# Multipole vectors!

Spherical Harmonics:

$$\frac{\delta T}{T}(\theta,\phi) = \sum_{l,m} a_{lm} Y_{lm}(\theta,\phi), \qquad C_\ell \equiv \frac{1}{2\ell+1}\sum_{m=-\ell}^{\ell} |a_{\ell m}|^2$$

Multipole Vectors:

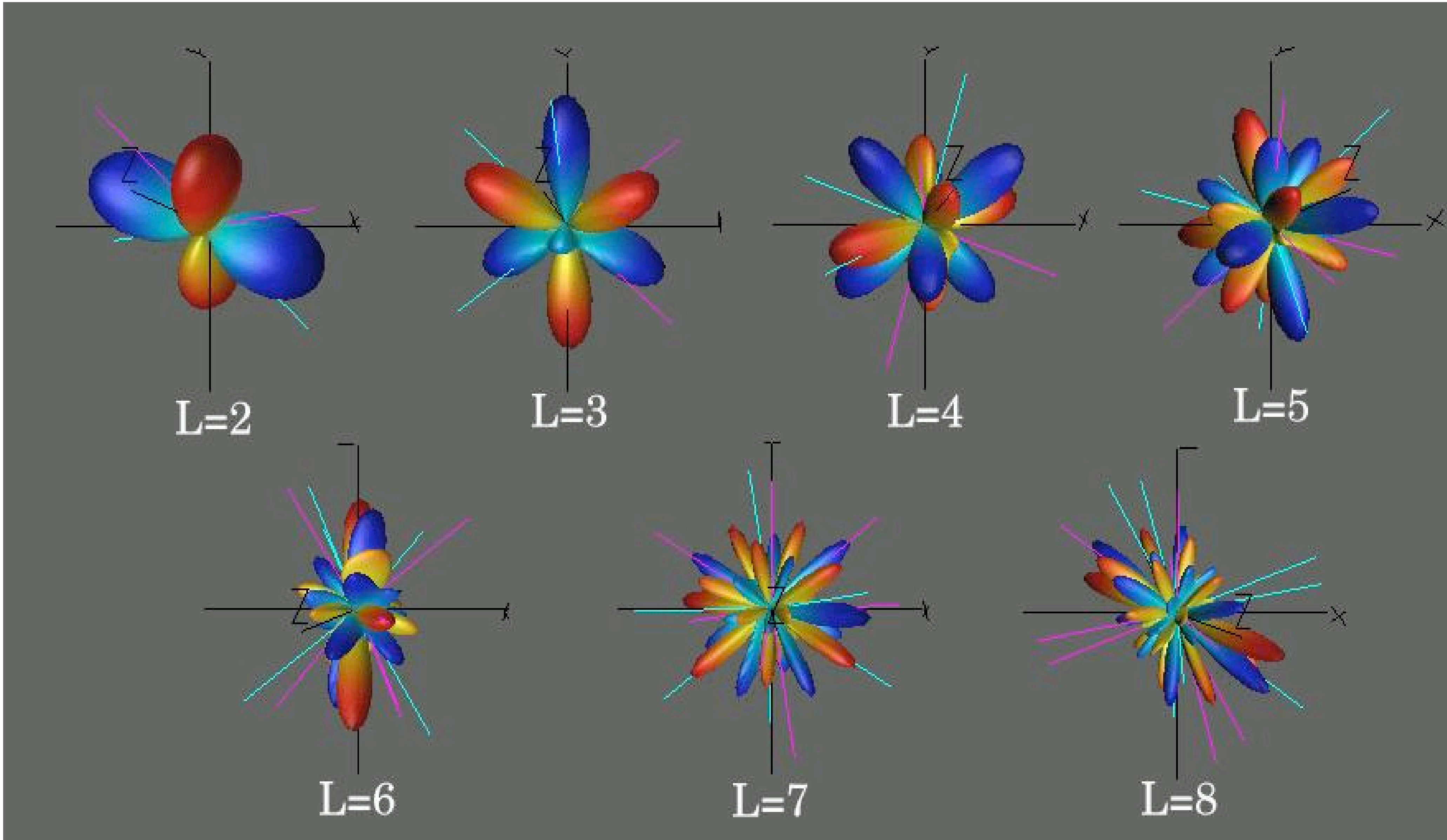
$$\sum_{m=-\ell}^{\ell} a_{lm} Y_{lm}(\theta,\phi) = A^{(\ell)} \left(\mathbf{v}_1^{(\ell)}\cdot\mathbf{e}\right)\cdots\left(\mathbf{v}_\ell^{(\ell)}\cdot\mathbf{e}\right)$$

$$\text{``}a^{(\ell)}_{i_1\ldots i_l} \leftrightarrow A^{(l)}\left[\mathbf{v}_1^{(\ell)}\otimes\mathbf{v}_2^{(\ell)}\otimes\ldots\mathbf{v}_\ell^{(\ell)}\right]\text{''}$$

Lth multipole <=> L (headless) vectors, plus a constant

Copi, Huterer & Starkman 2003; http://www.phys.cwru.edu/projects/mpvectors/

# Multipole vectors of our sky

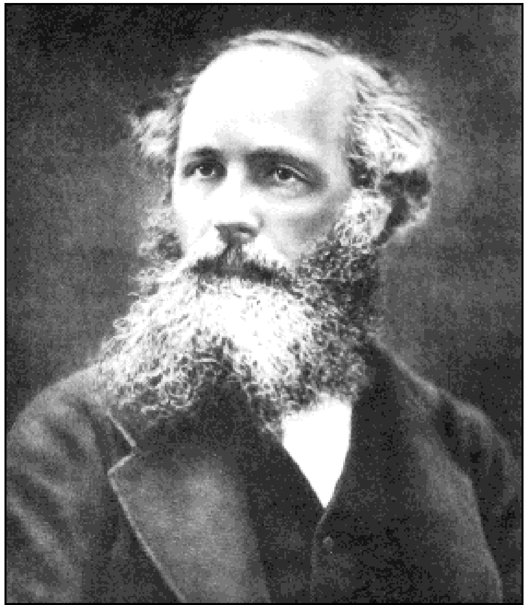



Copi, Huterer & Starkman 2003; http://www.phys.cwru.edu/projects/

# Why multipole vectors?

- A different representation of the CMB sky than the spherical harmonics, related highly non-linearly
- Well suited for looking for planarity/directionality
- Many interesting properties, theorems (Katz & Weeks 2004, Weeks 2005, Lachieze-Rey 2004, Dennis 2005...)
- (Reviewed in Copi, Huterer, Schwarz & Starkman astro-ph/0508047)

Also:
discussed by J.C. Maxwell in his
"Treatise on Electricity and Magnetism"
in 1892!!



# Maxwell's multipole vectors

Potential of:

Dipole: $\nabla_{\mathbf{v_1}} \frac{1}{r} \quad \left[= -\frac{\mathbf{v_1} \cdot \mathbf{r}}{r^3}\right]$

Quadrupole: $\nabla_{\mathbf{v_2}} \nabla_{\mathbf{v_1}} \frac{1}{r} \quad \left[= \frac{3(\mathbf{v_1} \cdot \mathbf{r})(\mathbf{v_2} \cdot \mathbf{r}) - r^2 (\mathbf{v_1} \cdot \mathbf{v_2})}{r^5}\right]$

..........

l'th multipole: $\nabla \mathbf{v}_\ell \ldots \nabla_{\mathbf{v_2}} \nabla_{\mathbf{v_1}} \frac{1}{r}$

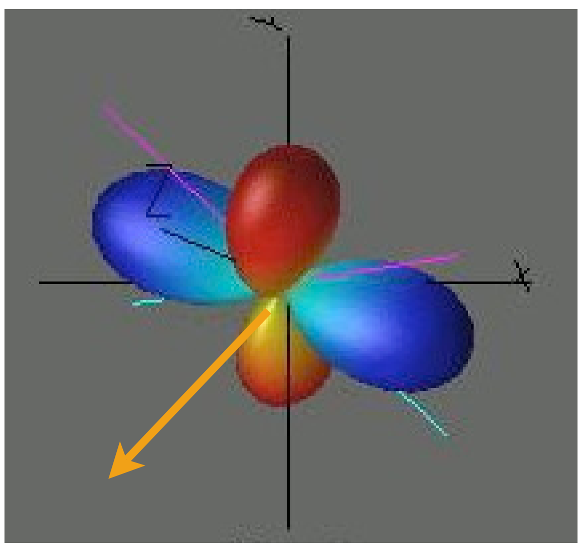
$\mathbf{v_1} \ldots \mathbf{v}_\ell$ are the multipole vectors

Maxwell 1892; Weeks 2004

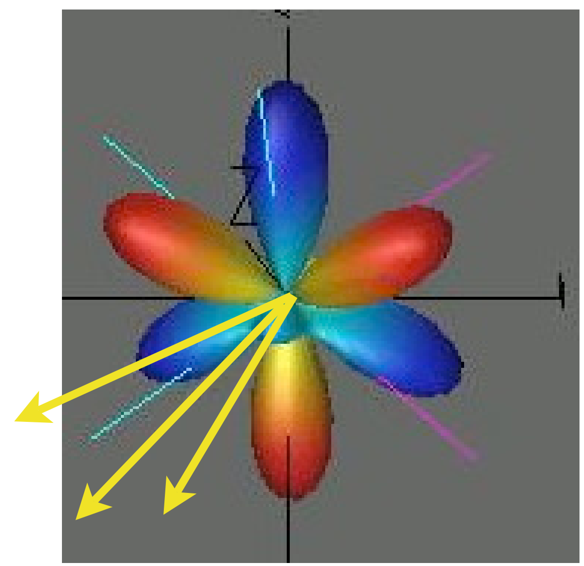
# Normals to multipole vectors

$$\mathbf{w}_{ij}^{(\ell)} \equiv \pm\left(\mathbf{v}_i^{(\ell)} \times \mathbf{v}_j^{(\ell)}\right)$$

"oriented areas"

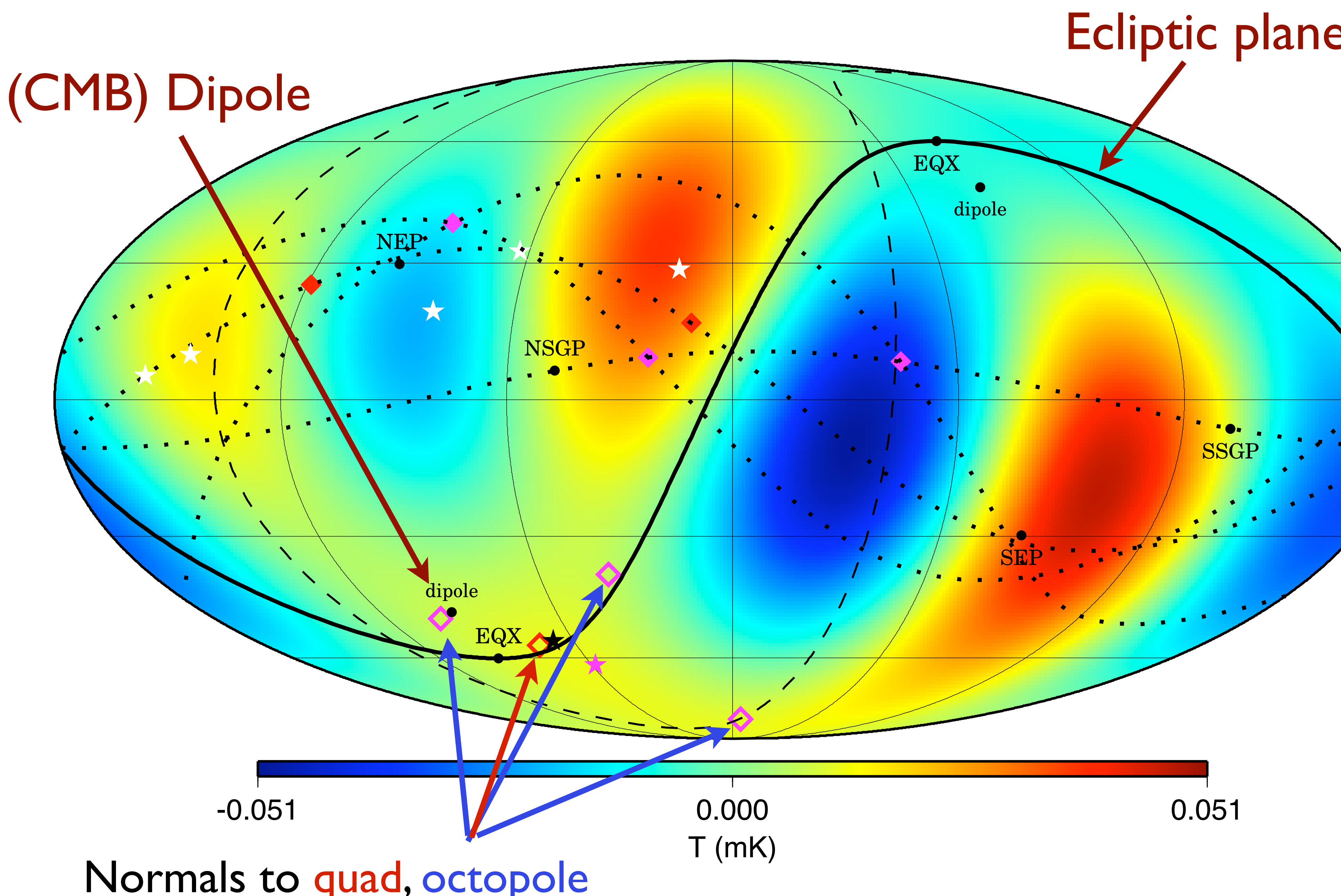


L=2



L=3

# L=2+3 alignments

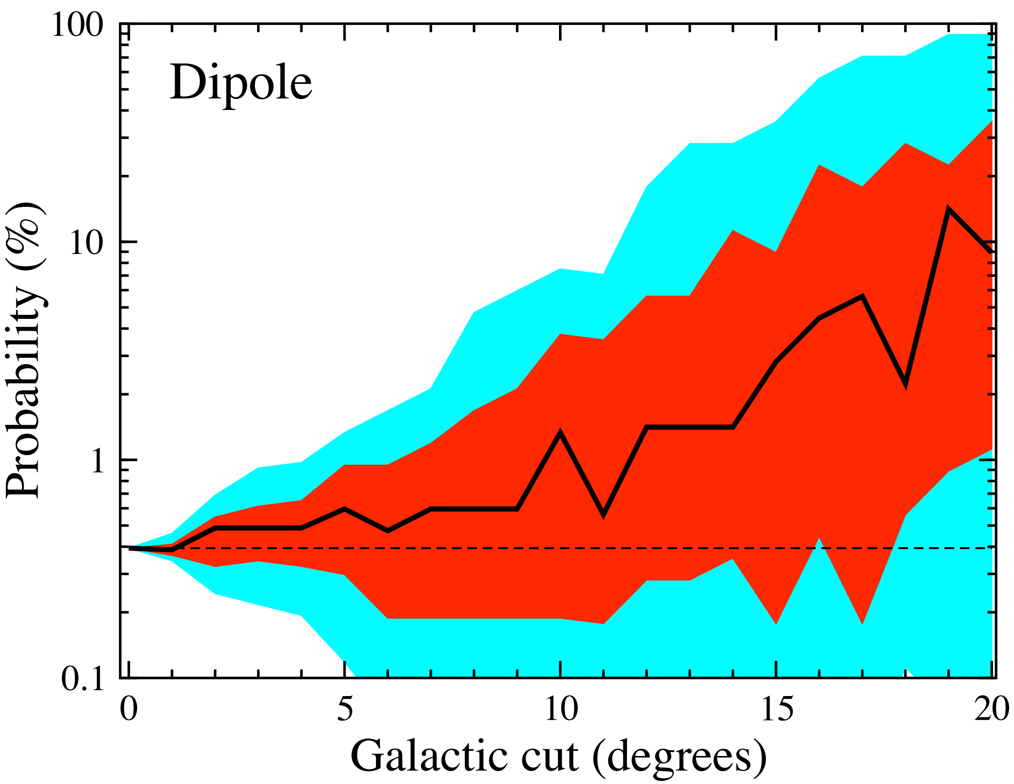


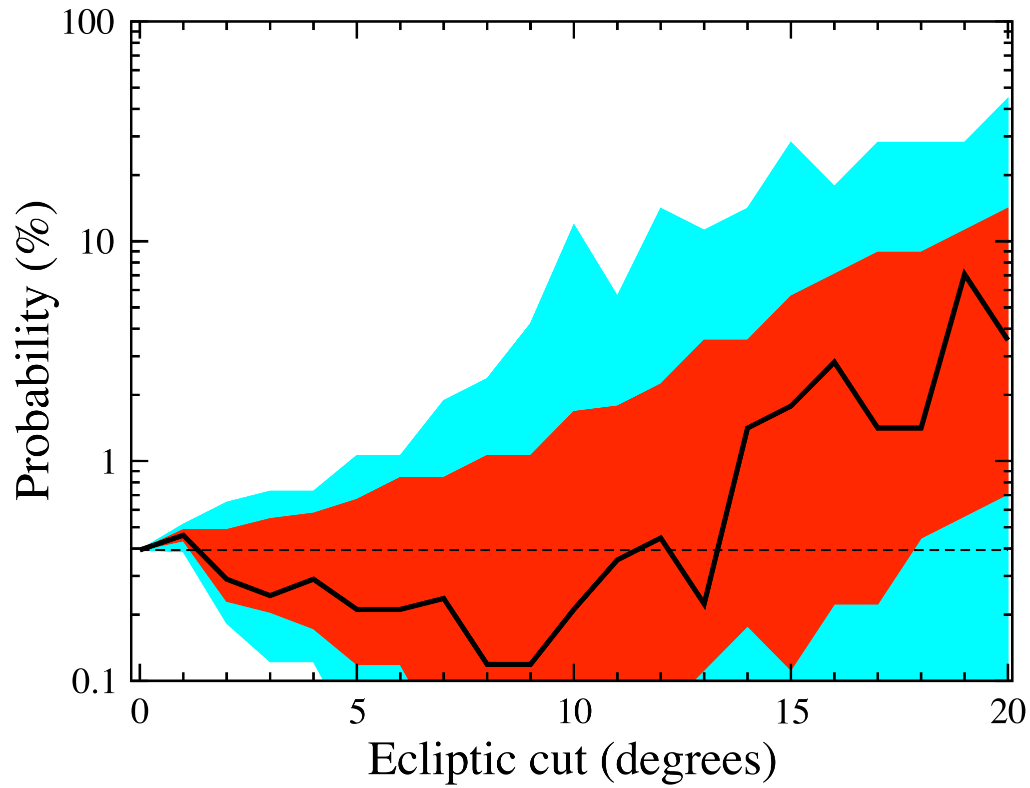

Schwarz, Starkman, Huterer & Copi, PRL, 2004

# Alignments found at L=2, 3

- The four area vectors are mutually close (99.0-99.9% CL)
- They lie close to ecliptic plane (98%-99% CL)
- They lie close to equinoxes and dipole (99.8% CL)
- Ecliptic plane carefully separates weak from strong extrema (93%-99.6% CL)

# Systematic checks: sky cut

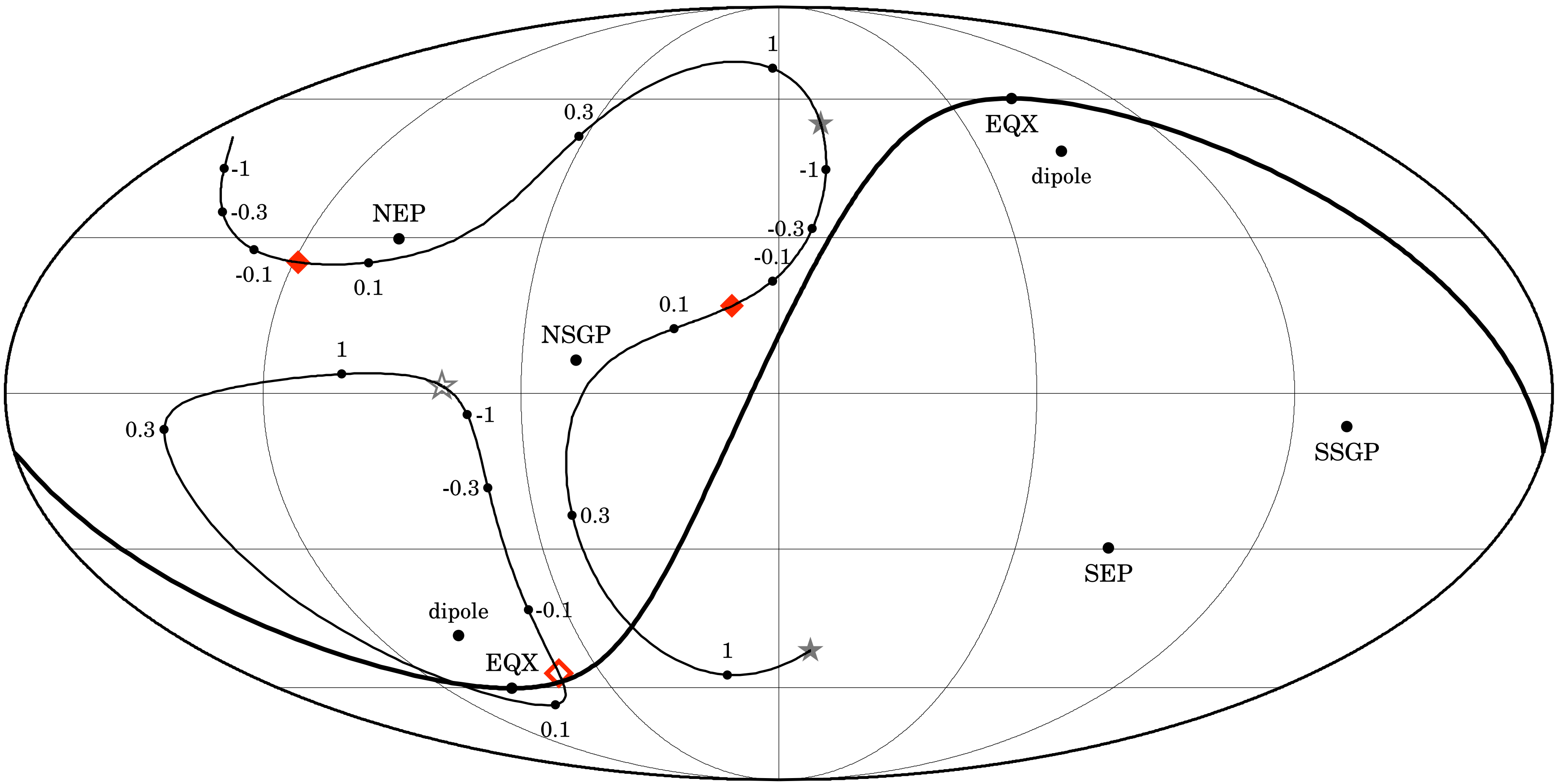



Errors increase sharply, but results consistent with full-sky result

Copi, Huterer, Schwarz & Starkman, MNRAS, 2006

# Systematic checks: foreground missubtraction

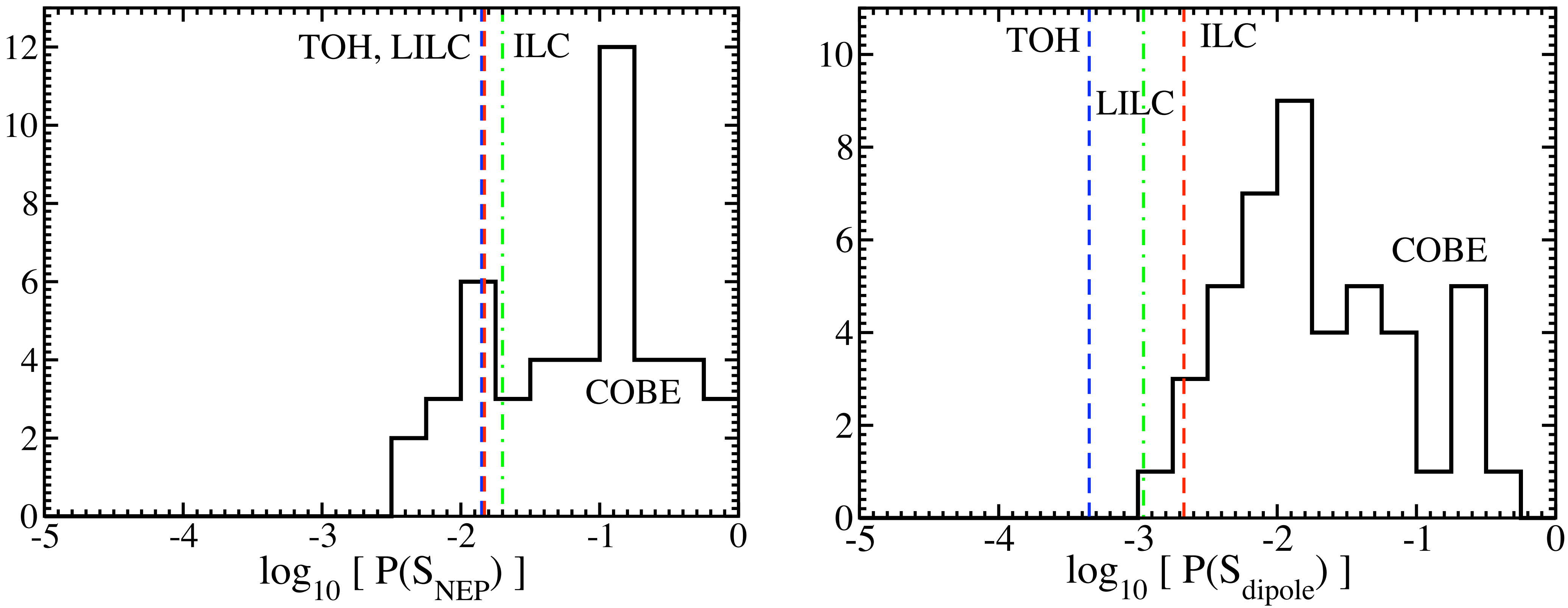



Adding (known) foregrounds leads to galactic, and not ecliptic, alignments

Copi, Huterer, Schwarz & Starkman, MNRAS, 2006

# What about COBE?

Using COBE MCMC maps from Wandelt, Larson & Lakshminarayanan 2003




Copi, Huterer, Schwarz & Starkman, MNRAS, 2006

# 4 classes of explanations:

- Astrophysical (e.g. an object or other source of radiation in the Solar System)
  - BUT: we think we know the Solar System. It would need to be a large source *and* undetected in data cross-checks.
- Instrumental (e.g. there is something wrong with WMAP instrument measuring CMB at large scales)
  - BUT: the instruments have been extremely well calibrated and checked. Plus, why would they pick out the Ecliptic plane?
- Cosmological (e.g. some property of the universe – inflation or dark energy for example – that we do not understand)
  - This is the most exciting possibility. BUT: why would the new/unknown physics pick out the Ecliptic plane?
- These alignments are a pure fluke!
  - BUT: they are <0.1% likely!

# Example: non-linear detector

Suppose that the WMAP detectors are slightly (1%) nonlinear

$$T_{\rm obs}(\hat{\mathbf{n}}) = T(\hat{\mathbf{n}}) + \alpha_2 T(\hat{\mathbf{n}})^2 + \alpha_3 T(\hat{\mathbf{n}})^3 + \ldots$$

The biggest signal on the sky is the dipole

$$T(\hat{\mathbf{n}}) = 3.3mK \cos(\theta)$$

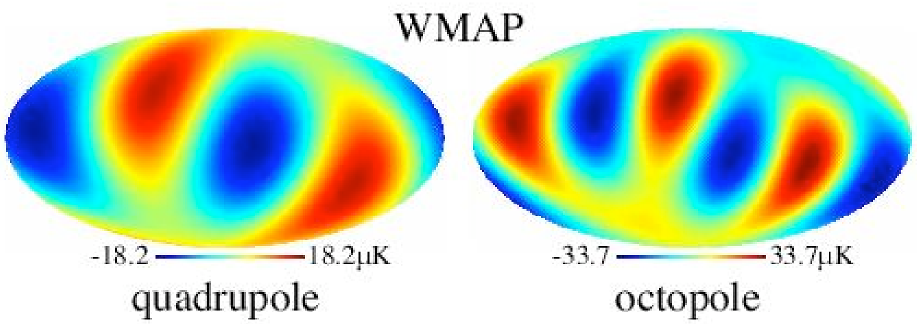
So with $\alpha_2 \sim \alpha_3 \sim 10^{-2}$, dipole anisotropy is modulated into a $10^{-5}$ quadrupole and octopole with $m = 0$ in the dipole frame.

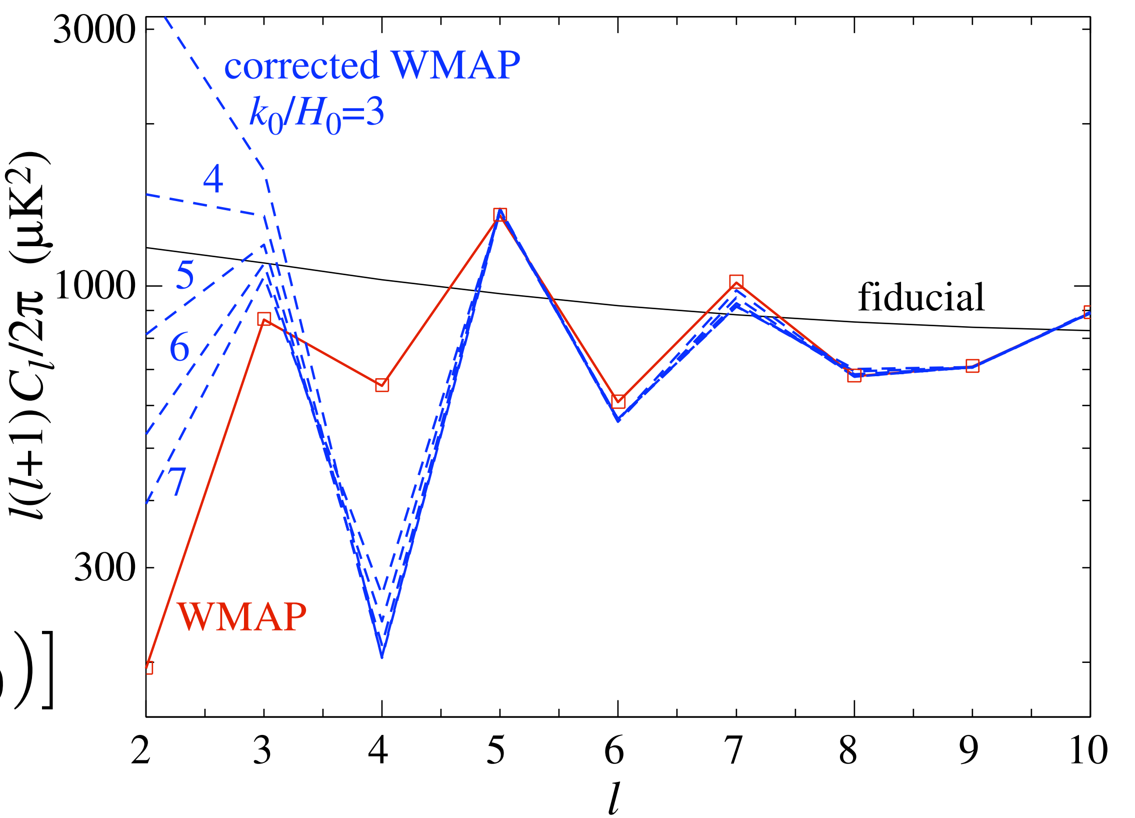
Sadly: doesn't work since would have been seen when observing $\sim 1K$ sources (in lab, Jupiter, etc).

Gordon, Hu, Huterer & Crawford, PRD, 2006

# Example: Spontaneous Isotropy Breaking

- To explain/model the apparent lack of isotropy on largest scales seen by WMAP




$$V(\phi) = V_0[1 + f\cos(\phi/M_0)]$$

$$\phi(z) = A + Bz$$

Modulates the CMB anisotropy through the ISW effect

Nonlinear modulation ⇨ a range of multipoles affected

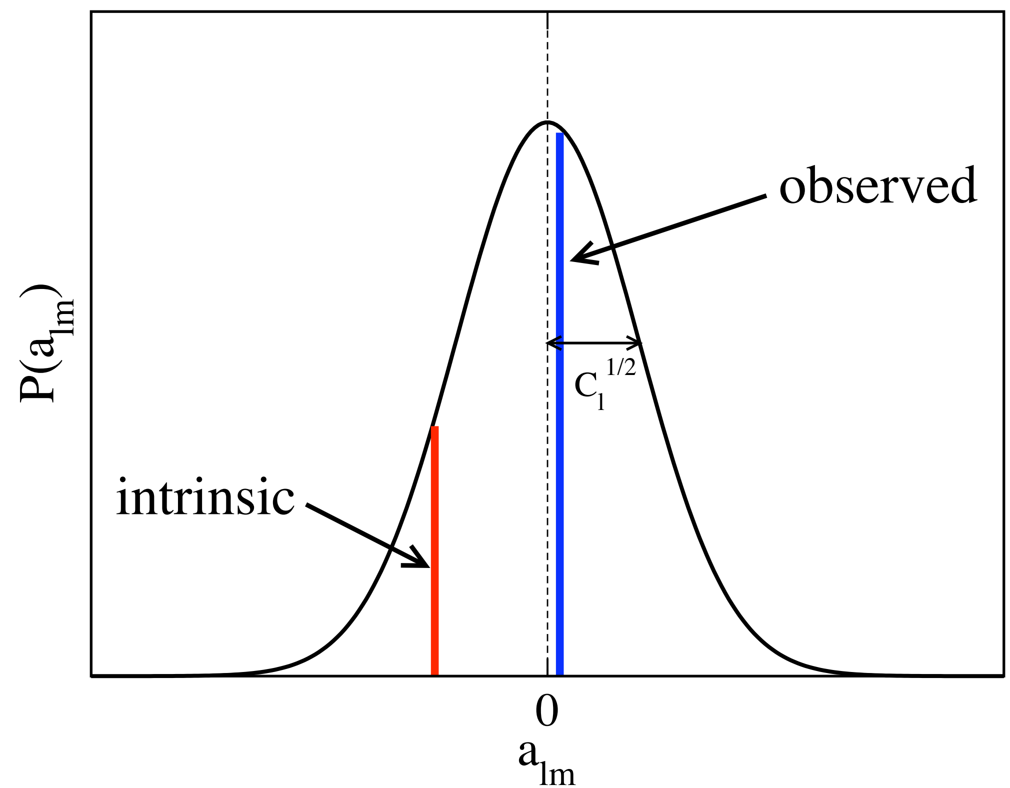
Gordon, Hu, Huterer & Crawford 2006

# Additive schemes “don’t work”

$$\hat{T}(\hat{\mathbf{n}}) = T_{\rm intr}(\hat{\mathbf{n}}) + T_{\rm extra}(\hat{\mathbf{n}})$$

Double (likelihood) penalty:

- Intrinsic sky is less likely than observed
- Requires a chance cancellation



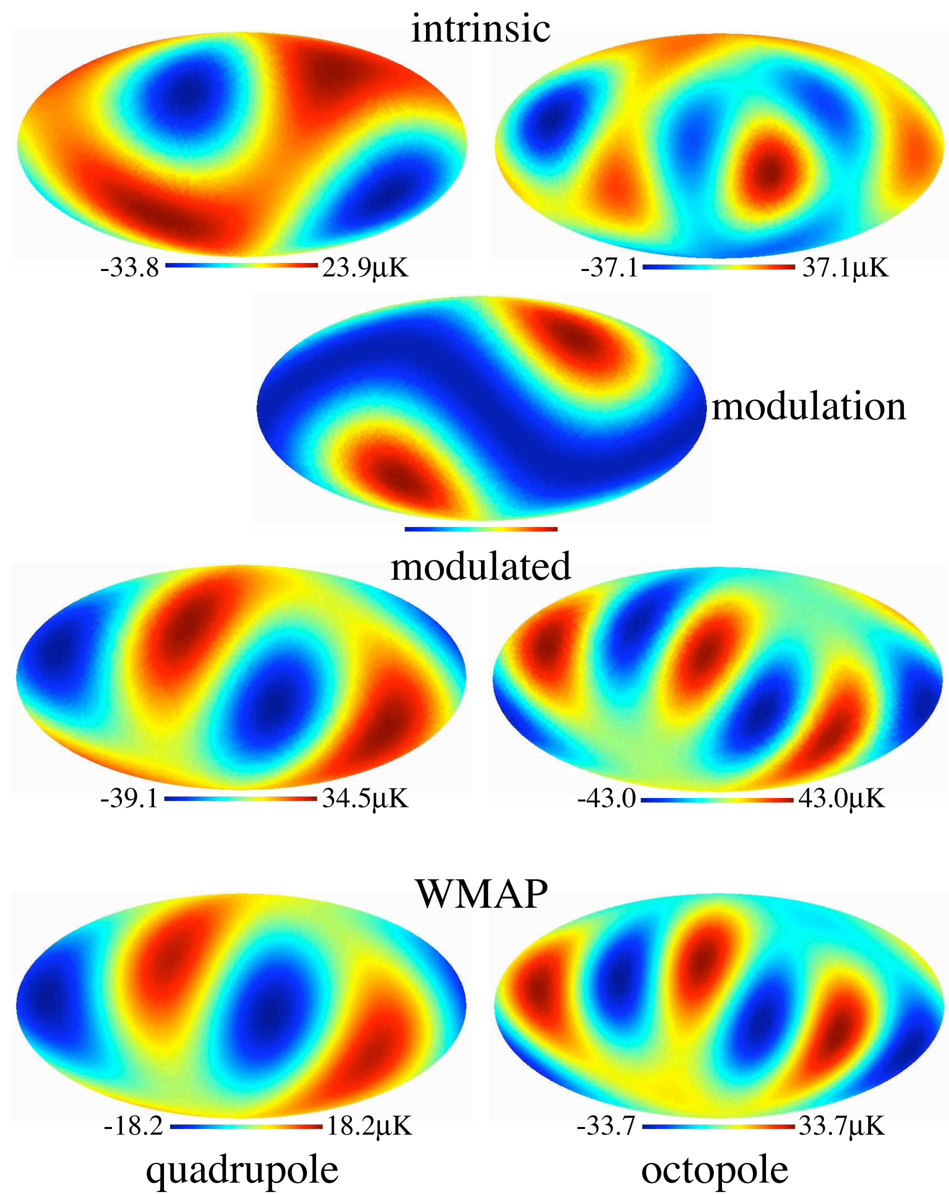

True for all additive schemes:
Solar System contamination,
Bianchi models,
etc

Gordon, Hu, Huterer & Crawford, PRD, 2006

# Multiplicative modulation can work

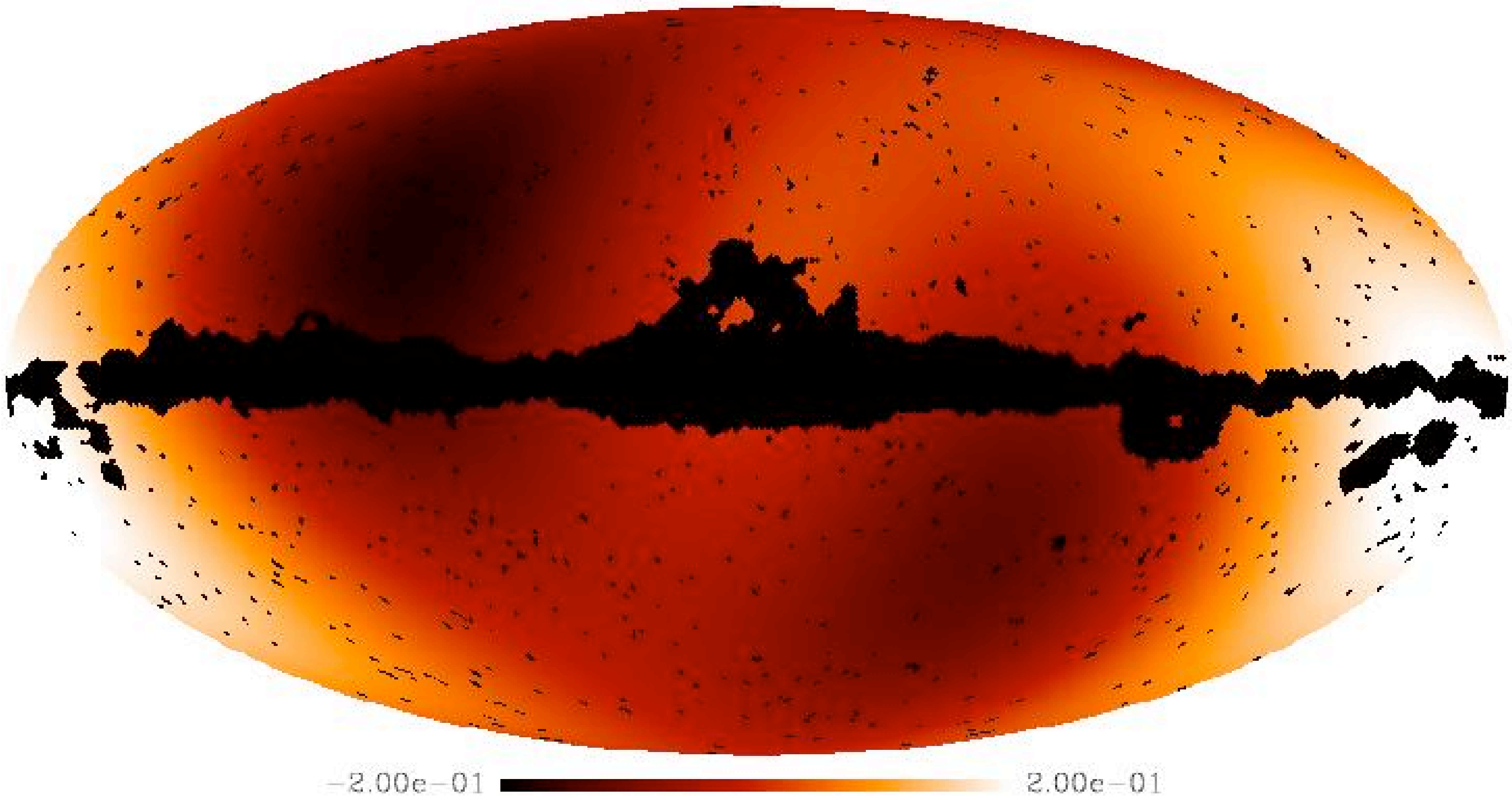
$$\hat{T}(\hat{\mathbf{n}}) = T_{\rm intr}(\hat{\mathbf{n}})\,[1 + w(\hat{\mathbf{n}})]$$

$w(\hat{\mathbf{n}}) \propto Y_{20}(\hat{\mathbf{n}})$ example



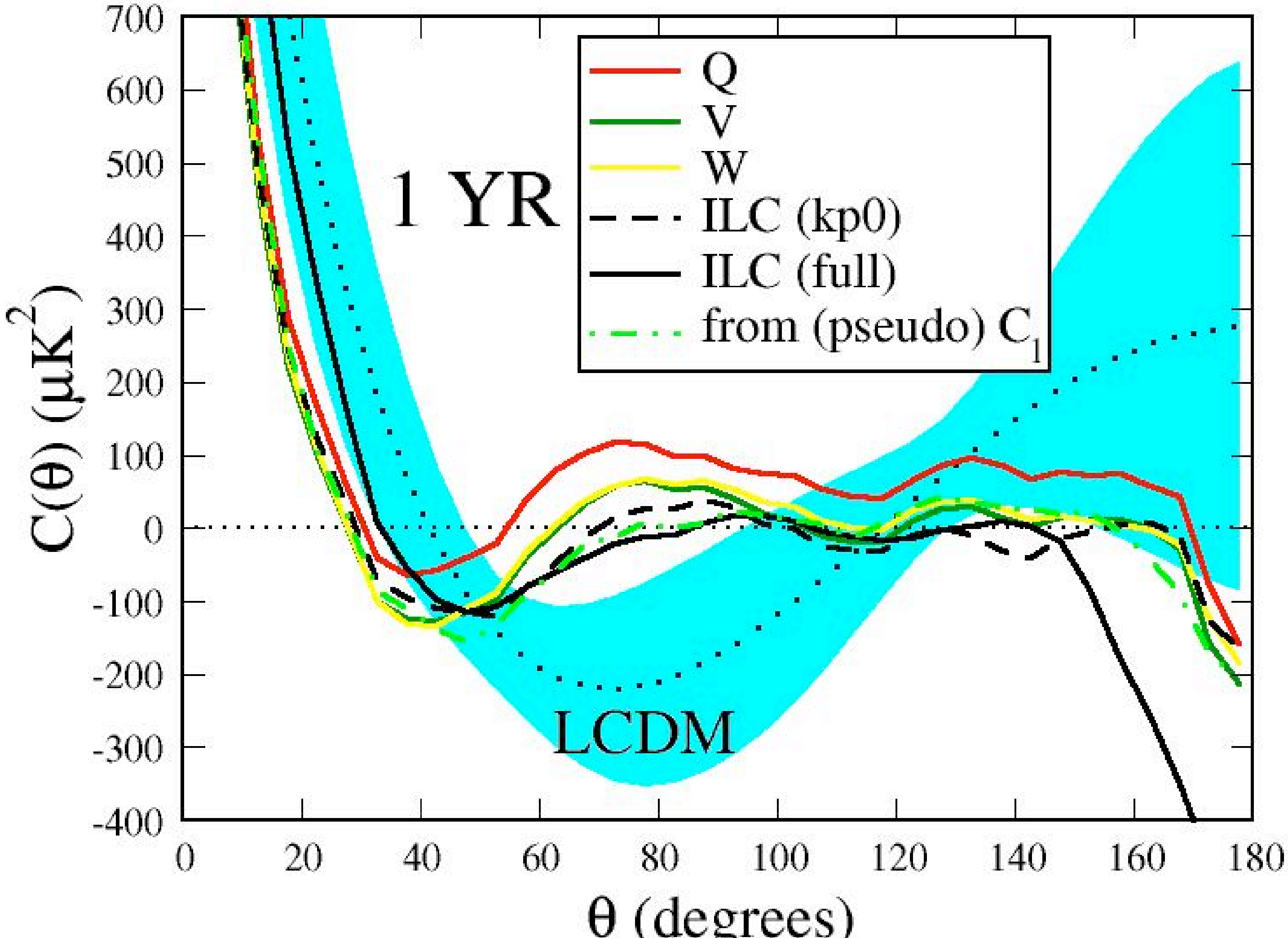

Gordon, Hu, Huterer & Crawford, PRD, 2006

# Best-fit L=1,2 multiplicative modulation from WMAP 123



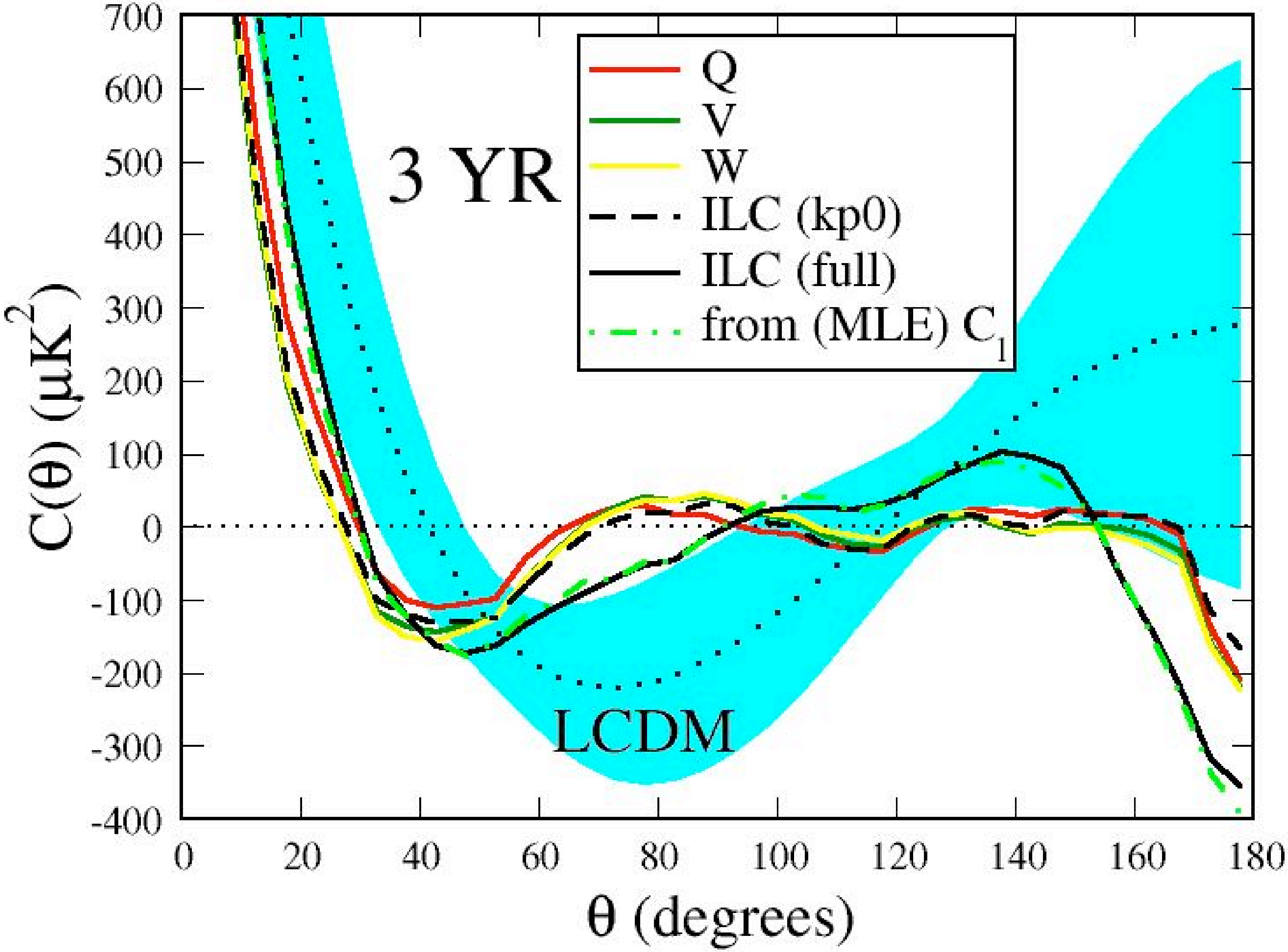

Spergel et al, 2006

# Low power on large scales

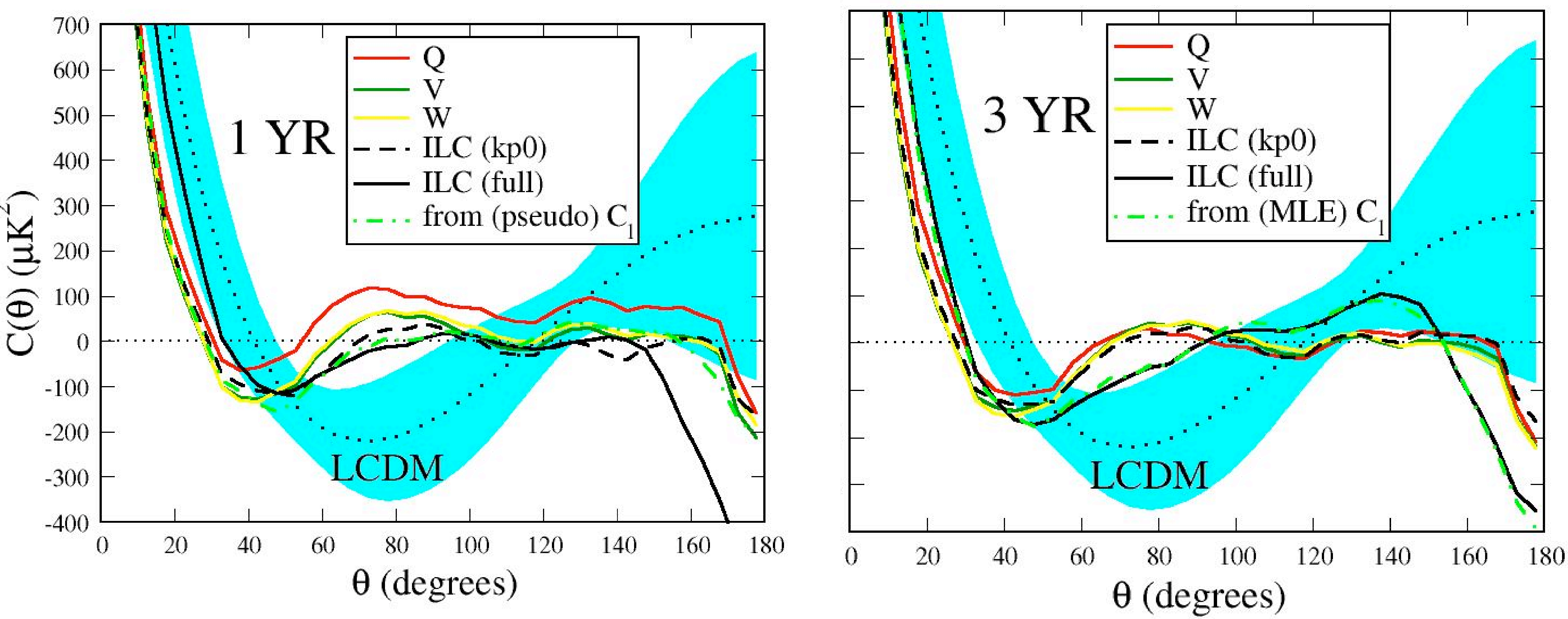



Spergel et al 2003: 0.2% of sims have less power at angles >60 deg

Copi, Huterer, Schwarz & Starkman astro-ph/0605135




Copi, Huterer, Schwarz & Starkman astro-ph/0605135




Copi, Huterer, Schwarz & Starkman astro-ph/0605135

# Future data and prospects

- WMAP is probably as good as it will get on large scales (as seen in year 1 vs year 123)
- Nevertheless, understanding of fine details is improving and is crucial.
- Planck will provide a great check of these measurements (very different experiment)
- Polarization maps with relatively high S/N, when eventually available, will provide even more leverage.
- The level of expected polarization "alignments" is model dependent
- In principle, can map out largest-scale fluctuations from wide-field, large-volume large-scale structure surveys (perhaps LSST)

# Conclusions

- Alignments with the ecliptic plane and/or dipole are sufficiently significant to be very interesting despite the a posteriori nature of these observations
- No convincing explanations so far
- Other observed anomalies (N/S asymmetry, L=4-6 etc) very intriguing and possibly related
- Multipole vectors are a great tool to study alignments and directionalities in the CMB
- Pixel-space C(theta) low at 99.97% CL - even more than in year 1