

Measuring the Copernican Principle

Wessel Valkenburg
Instituut-Lorentz, Leiden University
ITP, Heidelberg University

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Lemaître-Tolman-Bondi

$$ds^2 = -dt^2 + S^2(r, t)dr^2 + R^2(r, t)(d\theta^2 + \sin^2 \theta d\phi^2)$$

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$$\left(\frac{\dot{R}}{R}\right)^2 = \left(\frac{\dot{a}}{a}\right)^2 = H^2(r, t) = \frac{8\pi\tilde{M}^2}{3} \left[\frac{1}{a^3(r, t)} + \frac{3k(r)}{4\pi a^2(r, t)} + \frac{\Lambda}{8\pi\tilde{M}^2} \right],$$

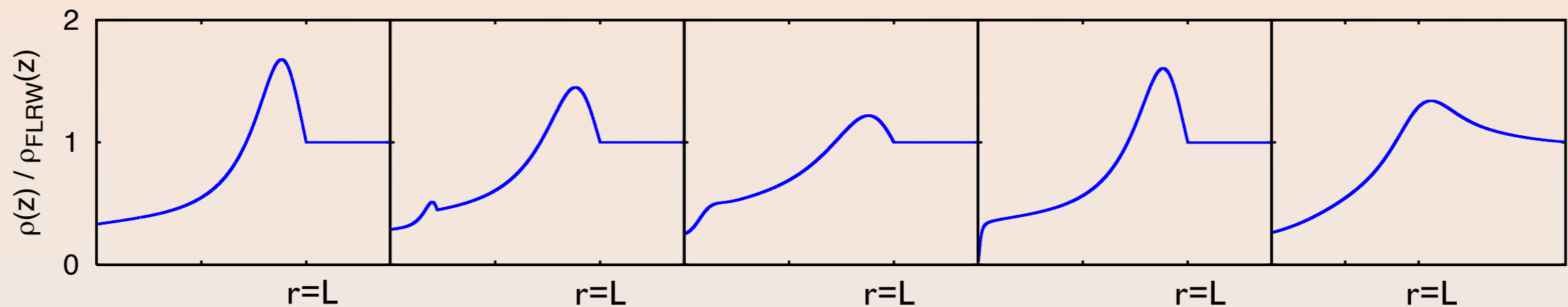
cosmological
constant

$$R(r, t) \equiv r a(r, t)$$

matter density

Radially dependent
curvature

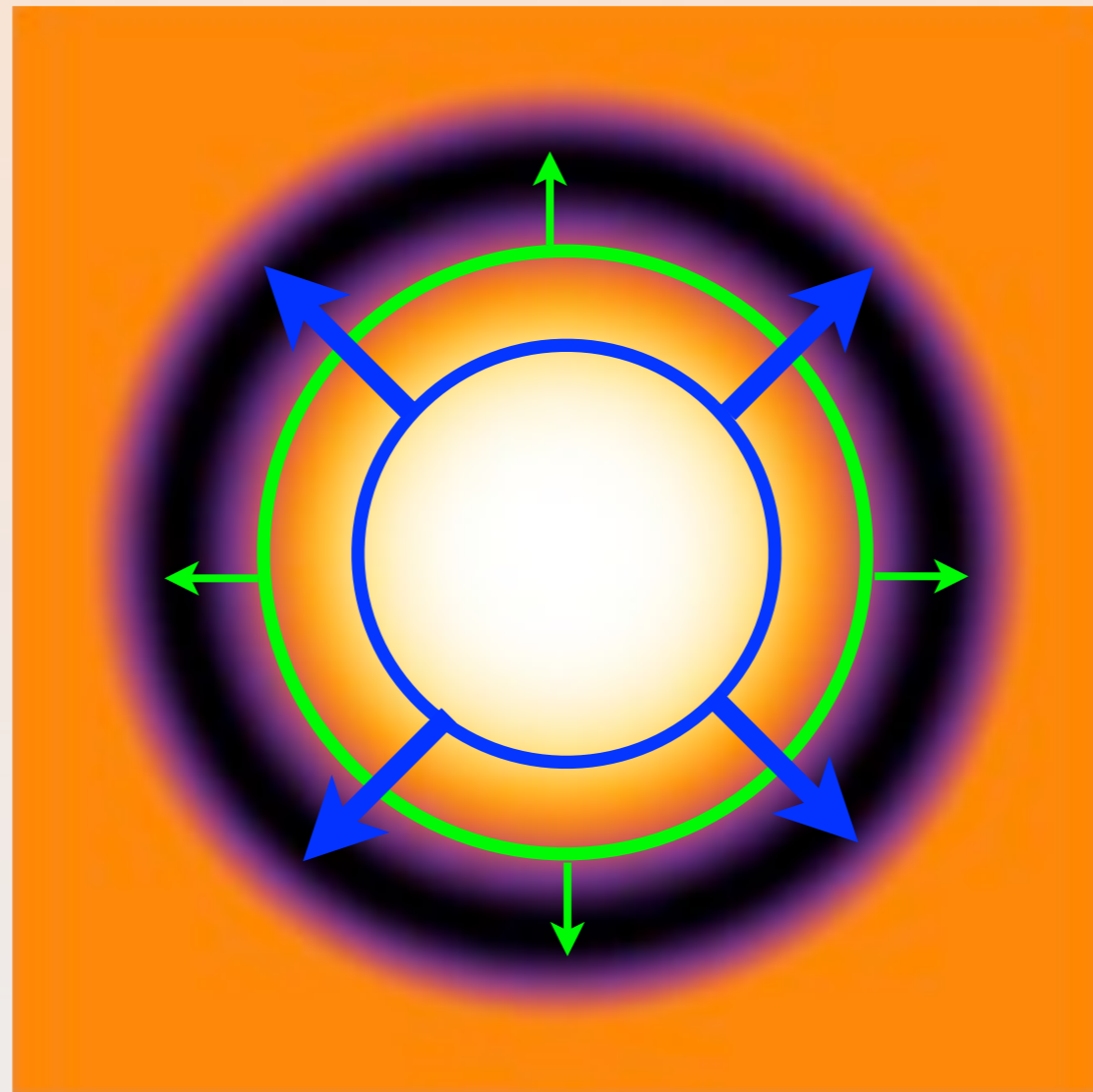
Once upon a time: no Λ



$$L \sim 1 - 2 \text{ Gpc}$$

Less matter \rightarrow more curvature \rightarrow slower decrease of H in time \rightarrow larger H at same time \rightarrow H in lightcone looks like accelerated expansion

Kinematic Sunyaev-Zeldovich effect



Kinematic Sunyaev-Zeldovich effect

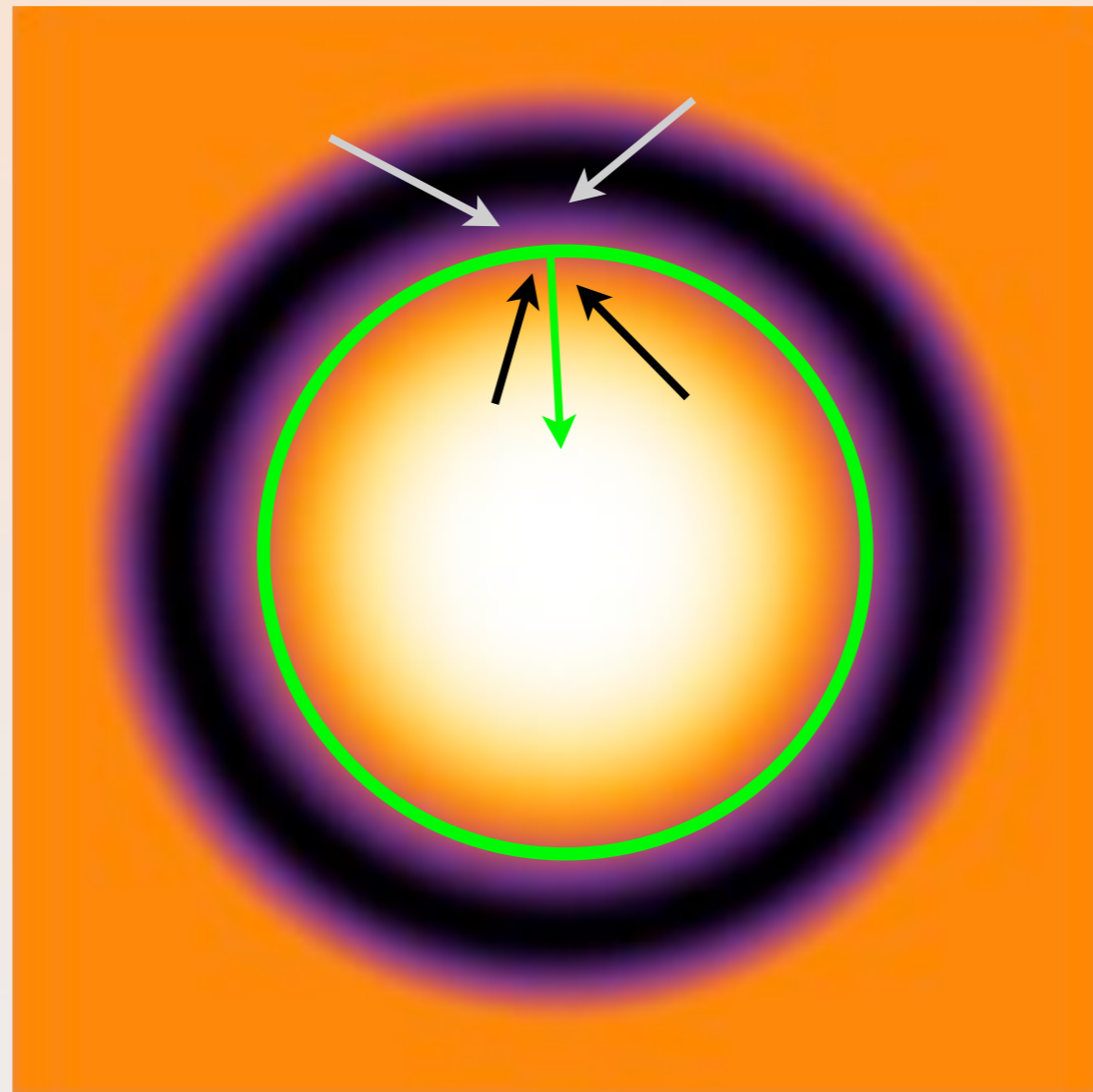
- Correlated outward velocity
- Transfer of energy from outward moving electron gas to CMB photons
- Strong imprint on CMB temperature map
- Rules out Gpc-scale Λ -less voids

[Garcia-Bellido, Haugbølle, 2008]

[Stebbins, Zhang, 2010]

[Zibin, Moss, 2011]

Compton-Y distortion



Black body spectrum to not entirely black body spectrum

Observational constraints on large perturbations around us

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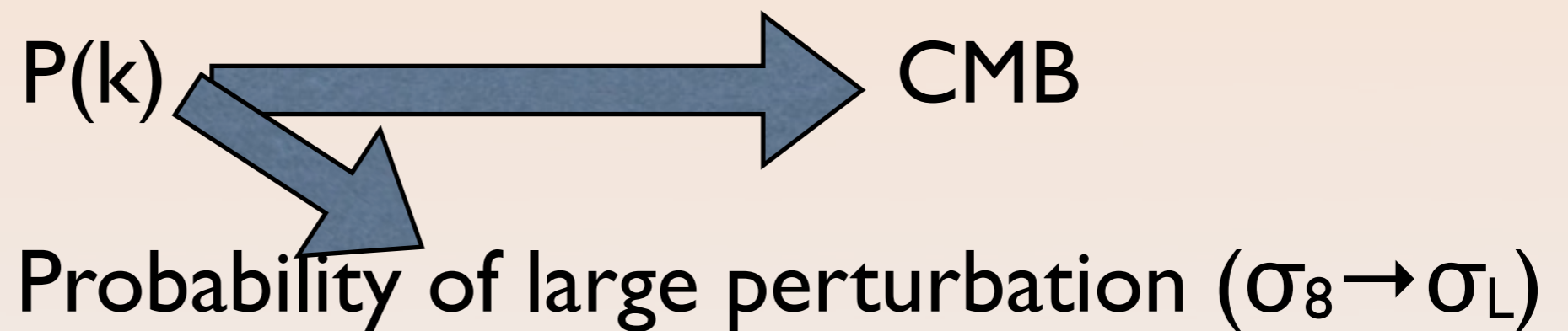
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- BAO: geometrical distance measure

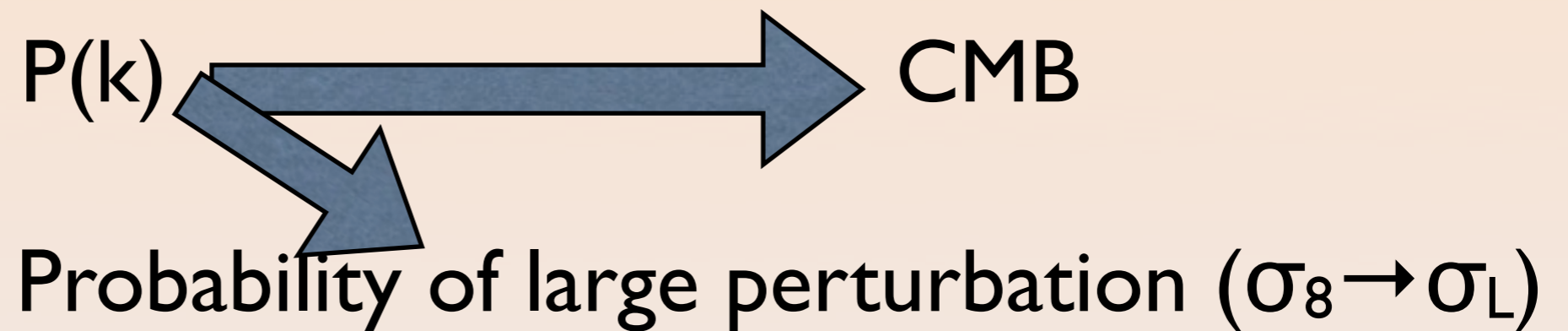
How copernican is the universe?

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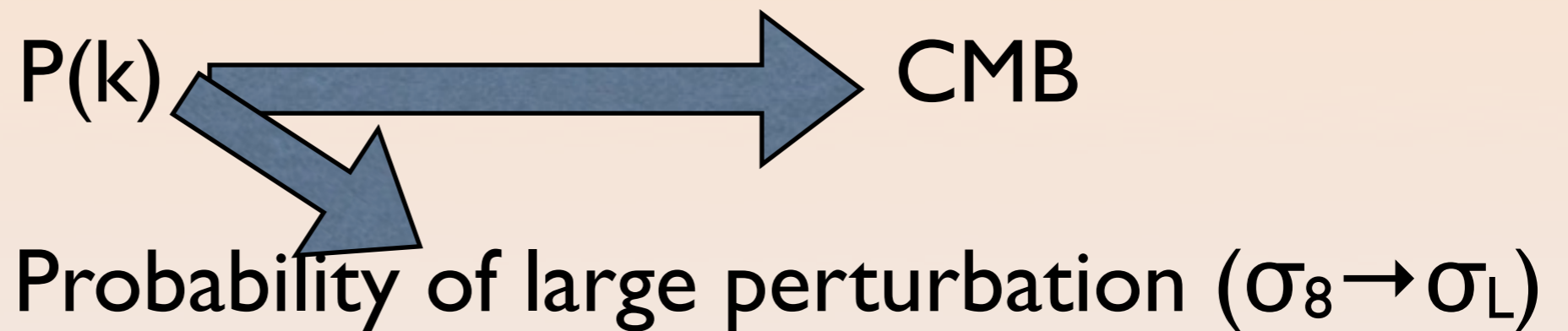
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Homogeneity Prior



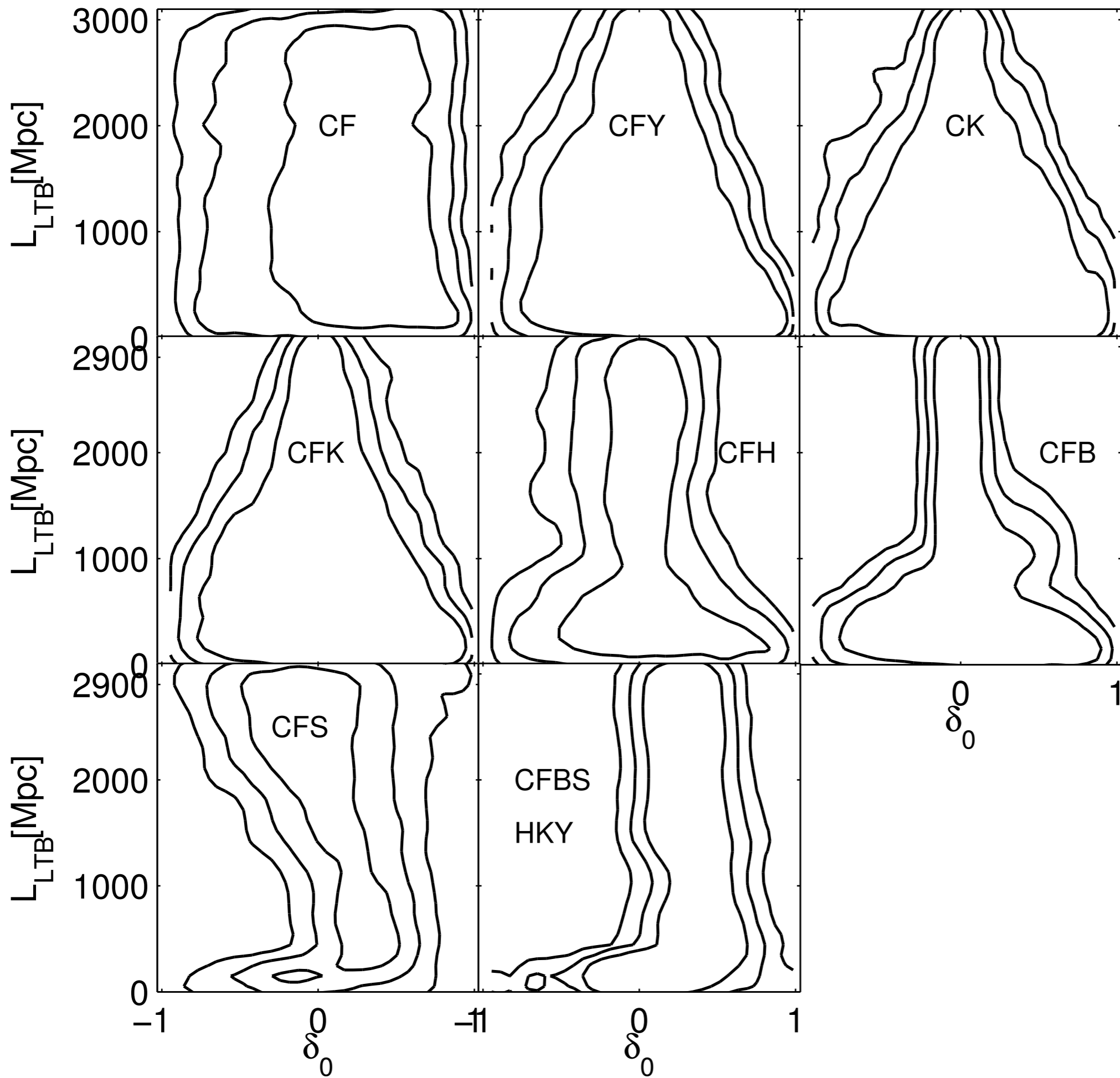
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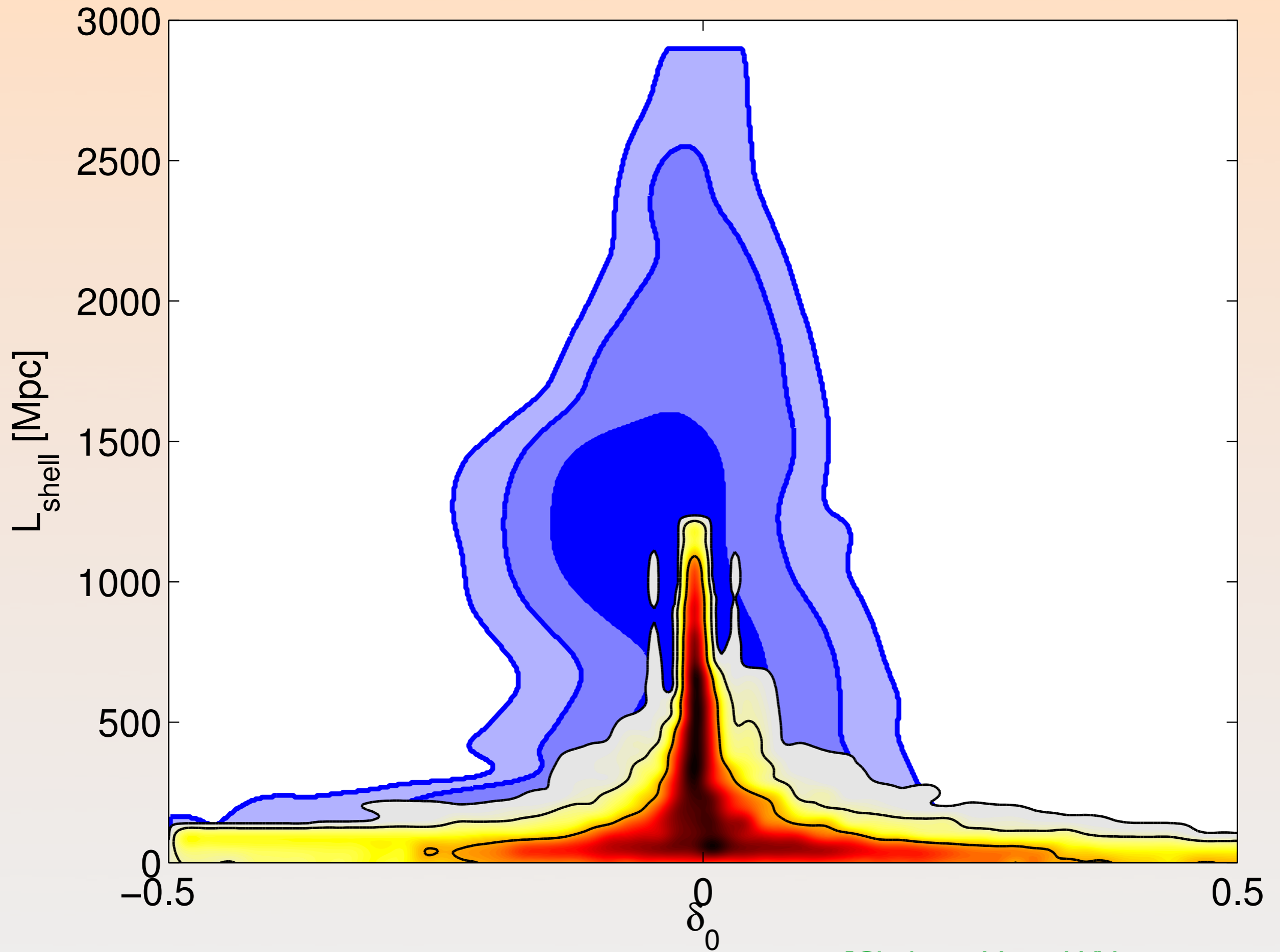


Observational constraints

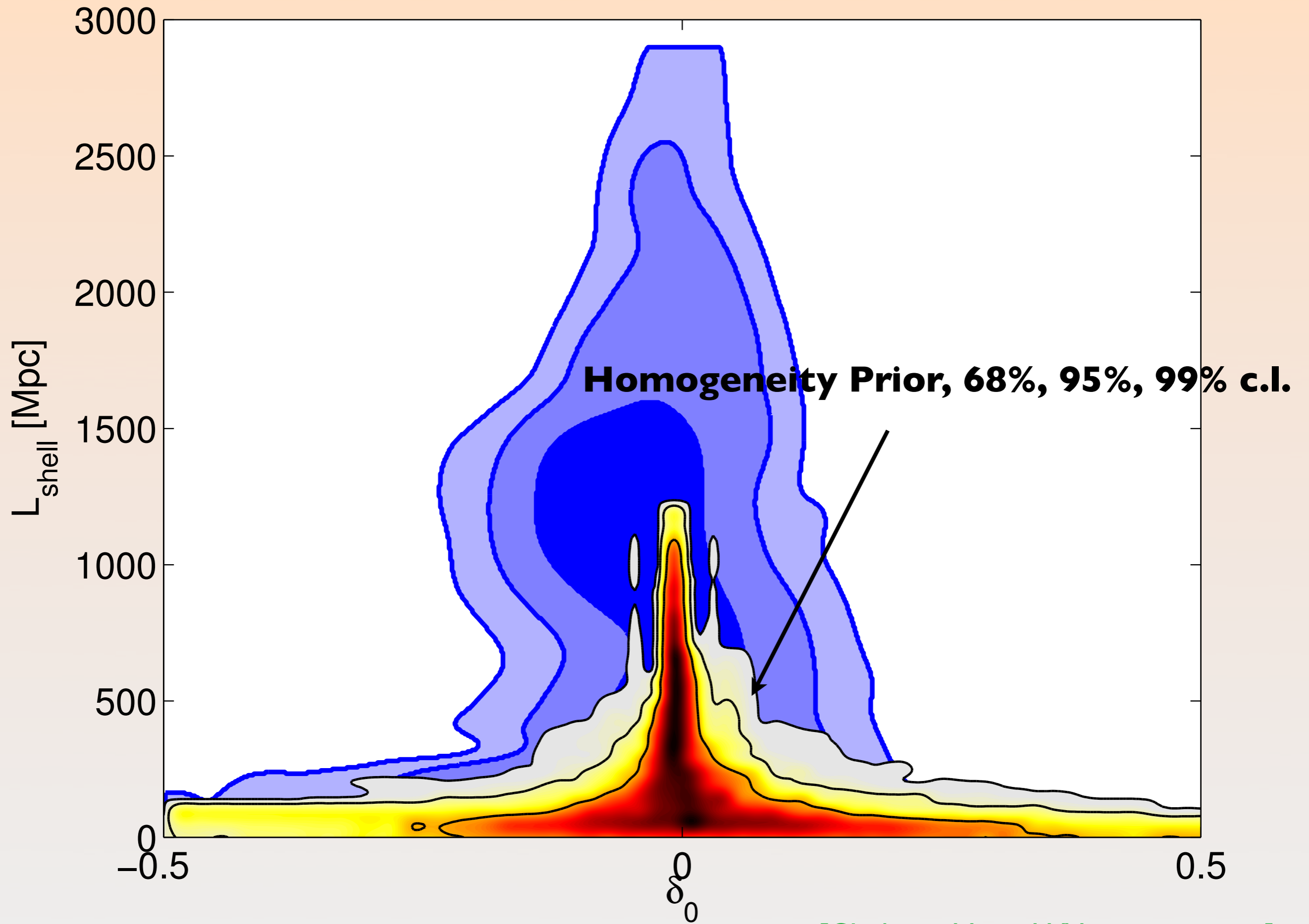
H_0 , SN, BAO, CMB, kSZ, Compton- y , $H(z)$



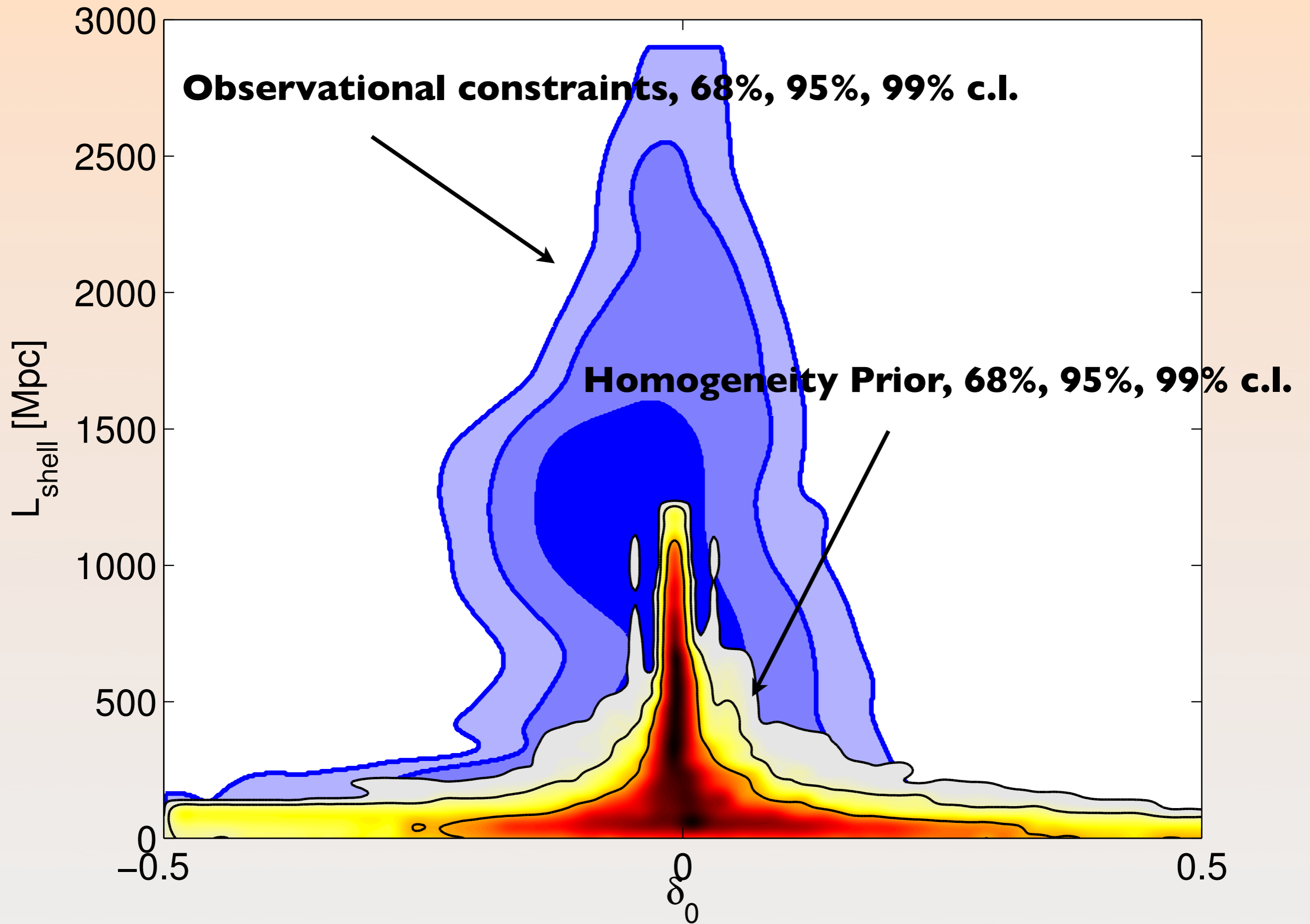
C=CMB
F= H_0
Y=Compton-Y
K=kSZ
H= $H(z)$
B=BAO
S=SN



[Clarkson, Marra, WV, in preparation]



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Summary

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Almost....